Dynamics of dark energy in collapsing halo of dark matter

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We investigate the non-linear evolution of spherical density and velocity perturbations of dark matter and dark energy in the expanding Universe. For this we have used the conservation and Einstein equations to describe the evolution of gravitationally coupled inhomogeneities of dark matter, dark energy and radiation from the linear stage in the early Universe to the non-linear stage at the current epoch. A simple method of numerical integration of the system of non-linear differential equations for evolution of the central part of halo is proposed. The results are presented for the halo of cluster ($k = 2\,\mathrm{Mpc}^{-1}$) and supercluster scales ($k = 0.2\,\mathrm{Mpc}^{-1}$) and show that a quintessential scalar field dark energy with a low value of effective speed of sound $c_s < 0.1$ can have a notable impact on the formation of large-scale structures in the expanding Universe.

Key words: dynamical dark energy, large scale structure of the Universe

INTRODUCTION

Dark energy is the mysterious dark component responsible for accelerated expansion of the Universe. It became an object of numerous studies in the past two decades. There are many possible explanations [1, 9] of the nature of dark energy. One of the most promising among them is scalar field dark energy, which can be modelled as a perfect fluid. In fact, with the knowledge of just a few parameters (the density $\rho_{\mathrm{DE}} = \rho_{cr}\Omega_{\mathrm{DE}}$, the equation of state (EoS) parameter $w_{\mathrm{DE}} = \rho/p$ and effective speed of sound $c_s^2 = \delta p/\delta\rho$), this model perfectly corresponds with the latest cosmological observational data [10, 13], while keeping a large number of possible parameter variations. Therefore, it would be interesting to study the behaviour of perturbations of such dark energy on non-cosmological scales. This may yield more constrains on parameters of the model.

It was already shown [2, 15] that non-linear perturbations of scalar field dark energy can in principle influence structure formation even on galaxy scales. Finally, in the past decade there was a number of papers (see [3] for review), which documented the studies of dark energy accretion in compact objects (black holes mainly) using the hydrodynamical approach. In our works [8, 14] we have shown that such stationary accretion does not change the gravitating mass of the central object, and in principle, it can notably affect the dynamics of bodies in its vicinity, which would be possible to extract from current observations.

Here we analyse the dynamics of dark energy in the collapsing halo of dark matter. In the next section we present the system of equations, which we use to solve the problem of evolution of spherical scalar perturbations in a 3-component expanding Universe from an early linear stage to a highly non-linear one, when a dark matter halo forms. In the last section we present numerical solutions and discuss their main features.

EQUATIONS FOR EVOLUTION OF PERTURBATIONS

We suppose that the Universe is spatially flat and filled by matter (m), dark energy (DE) and radiation (r), the metric of the background space-time is that of Friedmann-Robertson-Walker (FRW). Each component is described in the perfect fluid approximation by energy density ε, pressure p and four-velocity u^i. The equation of state for each component can be presented as $p = w\varepsilon$, with $w_{\mathrm{m}} = 0$ for matter, $w_{\mathrm{r}} = 1/3$ for radiation and $w_{\mathrm{DE}} < -1/3$ for dark energy. We assume that dark energy is a scalar field with $w_{\mathrm{DE}} = \mathrm{const}$. The goal of the paper is to analyse the evolution of a spherical halo from the linear stage in the early radiation-dominated epoch, through the turnaround point to the highly non-linear stage, infall of matter before virialization, at dark energy-dominated epoch. The local spherical perturbation distorts the FRW metric, so that it becomes

$$ds^2 = e^{\nu(t,r)}dt^2 - a^2(t)e^{\mu(t,r)}\times$$
$$\times \left[dr^2 + r^2 \left(d\theta^2 + \sin^2\theta d\varphi^2 \right) \right], \quad (1)$$

where the metric functions $\nu(t,r)$ and $\mu(t,r)$ vanish into the cosmological background. At the lin-

ear stage, when $\nu(t,r) \sim \mu(t,r) \ll 1$ the metric (1) becomes the metric of conformal-Newtonian (longitudinal) gauge [4] in spherical coordinates, since $e^{\nu(t,r)} \approx 1 + \nu$, $e^{\mu(t,r)} \approx 1 + \mu$. We rewrite the energy-momentum tensor of dark energy, $T^k_{i\,(\text{DE})} = (\varepsilon_{\text{DE}} + p_{\text{DE}})u_{i\,(\text{DE})}u^k_{\text{DE}} - \delta^k_i p_{\text{DE}}$, in terms of proper 3-velocity of fluid v_{DE} (measured in local frame), which only has a radial component. The relation between components of 4-velocity u^i_{DE} and 3-velocity v_{DE} of fluid are as follows

$$u_{i\,(\text{DE})} = \left\{ \frac{e^{\nu/2}}{\sqrt{1 - v^2_{\text{DE}}}}, -\frac{av_{\text{DE}}e^{\mu/2}}{\sqrt{1 - v^2_{\text{DE}}}}, 0, 0 \right\},$$

$$u^i_{\text{DE}} = \left\{ \frac{e^{-\nu/2}}{\sqrt{1 - v^2_{\text{DE}}}}, \frac{a^{-1}v_{\text{DE}}e^{-\mu/2}}{\sqrt{1 - v^2_{\text{DE}}}}, 0, 0 \right\},$$

where v_{DE} is in the units of speed of light. The non-zero components of energy-momentum tensor are

$$
\begin{aligned}
T^0_0 &= \frac{\varepsilon_{\text{DE}} + v^2_{\text{DE}}p_{\text{DE}}}{1 - v^2_{\text{DE}}}, \\
T^1_0 &= \frac{\varepsilon_{\text{DE}} + p_{\text{DE}}}{1 - v^2_{\text{DE}}}a^{-1}v_{\text{DE}}e^{(\nu-\mu)/2}, \\
T^1_1 &= -\frac{v^2_{\text{DE}}\varepsilon_{\text{DE}} + p_{\text{DE}}}{1 - v^2_{\text{DE}}}, \quad T^2_2 = T^3_3 = -p_{\text{DE}}.
\end{aligned}
$$

We decompose both density and pressure into background averaged and perturbed parts as $\varepsilon_{\text{DE}} = \bar{\varepsilon}_{\text{DE}}(1 + \delta_{\text{DE}})$, $p_{\text{DE}} = \bar{\varepsilon}_{\text{DE}}\left[w_{\text{DE}} + c^2_s\delta_{\text{DE}} - 3\dot{a}(1 + w_{\text{DE}})(c^2_s - w_{\text{DE}})\int e^{(\mu-\nu)/2}v_{\text{DE}}dr\right]$, where the density perturbation of each component is defined as $\delta(t,r) \equiv \left[\varepsilon(t,r) - \bar{\varepsilon}(t)\right]/\bar{\varepsilon}(t)$. We study the model of dark energy, for which both w_{DE} and c^2_s are constant. The integral term in the pressure decomposition comes from a non-adiabatic part of pressure perturbation of scalar field dark energy (details can be found in papers [5, 6, 11]). The presence of this term makes our equation of state non-barotropic, as we work in a frame different from that of proper dark energy and hence, the relation $p_{\text{DE}} = w_{\text{DE}}\rho_{\text{DE}}$ holds only for averaged parts of this component. Taking into account this decomposition, and keeping terms with v^0_{DE}, v^1_{DE} and v^2_{DE} only, we obtain the following energy-momentum tensor components of dark energy:

$$T^0_0 = \bar{\varepsilon}_{\text{DE}}(1 + \delta_{\text{DE}}) + \bar{\varepsilon}_{\text{DE}}(1 + w_{\text{DE}} + (1 + c^2_s)\delta_{\text{DE}})v^2_{\text{DE}},$$

$$T^1_0 = \bar{\varepsilon}_{\text{DE}}\left[1 + w_{\text{DE}} + (1 + c^2_s)\delta_{\text{DE}} - 3\dot{a}(1 + w_{\text{DE}})\times\right.$$
$$\left.\times(c^2_s - w_{\text{DE}})\int e^{(\mu-\nu)/2}v_{\text{DE}}dr\right]a^{-1}v_{\text{DE}}e^{(\nu-\mu)/2},$$

$$T^0_1 = -\bar{\varepsilon}_{\text{DE}}\left[1 + w_{\text{DE}} + (1 + c^2_s)\delta_{\text{DE}} - 3\dot{a}(1 + w_{\text{DE}})\times\right.$$
$$\left.\times(c^2_s - w_{\text{DE}})\int e^{(\mu-\nu)/2}v_{\text{DE}}dr\right]av_{\text{DE}}e^{(\mu-\nu)/2},$$

$$T^1_1 = -\bar{\varepsilon}_{\text{DE}}\left[w_{\text{DE}} + c^2_s\delta_{\text{DE}} - 3\dot{a}(1 + w_{\text{DE}})\times\right.$$
$$\left.\times(c^2_s - w_{\text{DE}})\int e^{(\mu-\nu)/2}v_{\text{DE}}dr\right] -$$
$$- \bar{\varepsilon}_{\text{DE}}\left[1 + w_{\text{DE}} + (1 + c^2_s)\delta_{\text{DE}}\right]v^2_{\text{DE}},$$

$$T^2_2 = T^3_3 = -\bar{\varepsilon}_{\text{DE}}\left[w_{\text{DE}} + c^2_s\delta_{\text{DE}} - 3\dot{a}(1 + w_{\text{DE}})\times\right.$$
$$\left.\times(c^2_s - w_{\text{DE}})\int e^{(\mu-\nu)/2}v_{\text{DE}}dr\right]. \quad (2)$$

Hereafter, one can get the corresponding equations and expressions for dark matter and radiation components from equations and expressions for dark energy just by putting $w_{\text{DE}} = c^2_s = 0$ for matter and $w_{\text{DE}} = c^2_s = 1/3$ for radiation.

To find the evolution of density and velocity perturbations we use two conservation equations, which have general covariant form

$$T^k_{i;k} = \frac{1}{\sqrt{-g}}\frac{\partial(\sqrt{-g}T^k_i)}{\partial x^k} - \frac{1}{2}\frac{\partial g_{kk}}{\partial x^i}g^{kk}T^k_k = 0$$

(for $i = 0$ it is continuity equation, for $i = 1$ it is motion equation). Substituting our energy-momentum tensor (2) one obtains:

$$\dot{\delta}_{\text{DE}}\left[1 + (1 + c^2_s)\delta_{\text{DE}}v^2\right] + 3\frac{\dot{a}}{a}\delta_{\text{DE}}(c^2_s - w_{\text{DE}}) +$$
$$+ \left[1 + w_{\text{DE}} + (1 + c^2_s)\delta_{\text{DE}}\right]\left[\frac{3\dot{\mu}}{2} + (1 - 3w_{\text{DE}} +\right.$$
$$\left.+ 2\dot{\mu})v^2_{\text{DE}} + 2\dot{v}_{\text{DE}}v_{\text{DE}}\right] + a^{-1}e^{(\nu-\mu)/2}\times$$
$$\times\left[(1 + c^2_s)\delta'_{\text{DE}}v_{\text{DE}} + \left\{1 + w_{\text{DE}} + (1 + c^2_s)\delta_{\text{DE}}\right\}\times\right.$$
$$\left.\times\left(v_{\text{DE}}(\nu' + \mu' + \frac{2}{r}) + v'_{\text{DE}}\right)\right] + 3\dot{a}(1 + w_{\text{DE}})\times$$
$$\times(c^2_s - w_{\text{DE}})\left[\left\{\frac{3\dot{a}}{a} + \frac{3\dot{\mu}}{2} + a^{-1}e^{(\nu-\mu)/2}\times\right.\right.$$
$$\left.\times\left(v_{\text{DE}}(\nu' + \mu' + \frac{2}{r}) + v'\right)\right\}\times$$
$$\left.\times\int e^{(\mu-\nu)/2}v_{\text{DE}}dr - v^2_{\text{DE}}\right] = 0, \quad (3)$$

$$\dot{v}_{\rm DE} + v_{\rm DE}\left(\frac{\dot{a}}{a}(1 - 3c_s^2) + 2\dot{\mu}\right) + \frac{\delta'_{\rm DE}e^{(\nu-\mu)/2}}{a(1 + w_{\rm DE})} \times$$

$$\times \left(c_s^2 + \frac{1 + c_s^2}{1 + w_{\rm DE}}v_{\rm DE}^2\right) + \frac{e^{(\nu-\mu)/2}}{a}\left(1 + \delta_{\rm DE}\frac{1 + c_s^2}{1 + w_{\rm DE}}\right) \times$$

$$\times \left[\frac{\nu'}{2} + (\nu' + \mu')v_{\rm DE}^2 + 2v_{\rm DE}\left(v'_{\rm DE} + \frac{v_{\rm DE}}{r}\right)\right] + \frac{1 + c_s^2}{1 + w_{\rm DE}} \times$$

$$\times \left[\dot{\delta}_{\rm DE}v + \dot{v}_{\rm DE}\delta + 2\dot{\mu}v_{\rm DE}\delta_{\rm DE} + (1 - 3w_{\rm DE})\frac{\dot{a}}{a}v_{\rm DE}\delta_{\rm DE}\right] -$$

$$-3\frac{\dot{a}}{a}(c_s^2 - w_{\rm DE})\left[\left(\dot{v}_{\rm DE} + (1 - 3w_{\rm DE})\frac{\dot{a}}{a}v_{\rm DE} + 2\dot{\mu}v_{\rm DE} + \right.\right.$$

$$\left. + \frac{\nu'}{2a}e^{(\nu-\mu)/2}\right)\int e^{(\nu-\mu)/2}v_{\rm DE}dr +$$

$$+ v_{\rm DE}\int e^{(\nu-\mu)/2}\left(\dot{v}_{\rm DE} + \frac{\ddot{a}}{a}v_{\rm DE} + \frac{\dot{\mu} - \dot{\nu}}{2}v_{\rm DE}\right)dr\right] = 0.$$
(4)

The continuity equation for background density, which we also use, is $\dot{\bar{\varepsilon}}_{\rm DE} + 3\frac{\dot{a}}{a}(1 + w_{\rm DE})\bar{\varepsilon}_{\rm DE} = 0$.

To find the metric functions we exploit the Einstein equations

$$R_j^i - \frac{1}{2}\delta_j^i R = \varkappa\left(T_j^i{}_{\rm (DE)} + T_j^i{}_{(m)} + T_j^i{}_{(r)}\right).$$

If we construct the equation $G_1^1 - G_2^2 = \varkappa\left(T_1^1 - T_2^2\right)$ then it becomes apparent that at the linear stages $\mu = -\nu$. At the non-linear stage, the right-hand side of this equation equals zero at the centre of perturbation, which again gives a reduction to only one potential. In this paper we analyse the dynamics of dark matter and dark energy in the central part of spherical overdensity only, therefore we accept $\mu = -\nu$ approximation, which gives us the possibility to use only one $\binom{0}{0}$ Einstein equation for determination of one metric function $\nu(t, r)$:

$$-3\frac{\dot{a}}{a}\dot{\nu} + \frac{3\dot{\nu}^2}{4} + 3\frac{\dot{a}^2}{a^2}\left(1 - e^\nu\right) + \frac{1}{a^2}\left[\nu'' + \frac{2}{r}\nu' + \frac{\nu'^2}{4}\right] =$$

$$= 3\frac{\dot{a}^2}{a^2}e^\nu\frac{\Omega_{\rm r}a^{-1}\delta_{\rm r} + \Omega_{\rm DE}a^{-3w_{\rm DE}}\delta_{\rm DE} + \Omega_{\rm m}\delta_{\rm m}}{\Omega_{\rm m} + \Omega_{\rm r}a^{-1} + \Omega_{\rm DE}a^{-3w_{\rm DE}}}. \quad (5)$$

Here we used the notations:

$$\Omega_N \equiv \bar{\varepsilon}_N^0 / \left(\bar{\varepsilon}_{\rm m}^0 + \bar{\varepsilon}_{\rm DE}^0 + \bar{\varepsilon}_{\rm r}^0\right),$$

where N denotes the type of fluid and "0" marks the value at current epoch. In the flat three component universe $\Omega_{\rm m} + \Omega_{\rm r} + \Omega_{\rm DE} = 1$. Eqs. (3)–(5) have non-relativistic limit (Appendix), which in case of dark matter ($c_s^2 = w_{\rm DE} = 0$) coincide with well known classical hydrodynamic equations and Poisson equation accordingly.

We are interested in the cluster and supercluster scales of perturbations, for which $\nu \ll 1$, $v_{\rm m} \ll 1$ in the bulk of object, while $\delta_{\rm m} \gg 1$ in their halos.

So, the equations can be essentially reduced by neglecting the terms like, ν^2, $\nu v_{\rm m}$, $\nu v_{\rm DE}$ and higher order terms. The terms with $v\nabla v$ must be kept, since they are important during the highly non-linear stage. The final reduced form of these equations, which we use in the code, can be found in our upcoming paper [11].

Therefore, we have seven 1st-order partial differential equations for seven unknown functions

$$\delta_{\rm m}(t, r), \quad v_{\rm m}(t, r), \quad \delta_{\rm DE}(t, r), \quad v_{\rm DE}(t, r),$$

$$\delta_{\rm r}(t, r), \quad v_{\rm r}(t, r), \quad \nu(t, r), \quad (6)$$

which can be solved numerically for given initial conditions.

At the early epoch the amplitudes of cosmological perturbations of space-time metric, densities and velocities are low and the Eqs. (3)–(5) can be linearized for all components. Moreover, we can present each function of (t, r) as a product of its amplitude, which depends on t only, and some function of radial coordinate r, which describes the initial profile of spherical perturbation, which can be expanded into series of some orthogonal functions, e. g. spherical ones in our case. In particular, we can present the perturbations of the metric, density and velocity of N-component as follows

$$\nu(t, r) = \tilde{\nu}(t)\frac{\sin kr}{kr}, \quad \delta_N(t, r) = \tilde{\delta}_N(t)\frac{\sin kr}{kr},$$

$$v_N(t, r) = \tilde{v}_N(t)\left(\frac{\sin kr}{kr}\right)' = \tilde{v}_N(t)k\left(\frac{\cos kr}{kr} - \frac{\sin kr}{k^2r^2}\right).$$

In the analysis of the evolution of the central part of a spherical halo we can decompose r-function in the Taylor series and only keep the leading terms:

$$f_k(r) \approx 1, \quad f'_k(r) \approx -\frac{1}{3}k^2r,$$

$$\int f'(r)dr \approx 1, \quad f''_k(r) + \frac{2f_k(r)}{r} \approx -k^2,$$

where $f_k(r) = \sin kr/kr$. It gives the possibility to reduce the system of seven partial differential equations for unknown functions (6) to the system of seven ordinary differential equations for their amplitudes

$$\tilde{\delta}_{\rm m}(t), \quad \tilde{v}_{\rm m}(t), \quad \tilde{\delta}_{\rm DE}(t), \quad \tilde{v}_{\rm DE}(t),$$

$$\tilde{\delta}_{\rm r}(t), \quad \tilde{v}_{\rm r}(t), \quad \tilde{\nu}(t),$$

which is presented in [11].

We set the adiabatic initial conditions in the following way. To find the relations between amplitudes at some $a_{\rm init} \ll 1$ when the scales of gravitationally bound systems were essentially larger than the horizon scale, $a_{\rm init}k^{-1} \gg t$ we used the linearization of system (3)–(5), which has an exact analytical solution for a one-component Universe. Evidently for cluster and supercluster scales the epoch

was the early radiation-dominated one when $\bar{\varepsilon}_\mathrm{r} \gg \bar{\varepsilon}_\mathrm{m} \gg \bar{\varepsilon}_\mathrm{DE}$, thus, the matter and dark energy can be treated as test components. The amplitude of the metric function $\tilde{\nu}$ is defined by density perturbations of the relativistic component. The non-singular solution of the corresponding equations has asymptotic values at $a_\mathrm{init}\tilde{\nu}^\mathrm{init} = -C$, $\tilde{\delta}_\mathrm{r}^\mathrm{init} = C$, $\tilde{v}_\mathrm{r}^\mathrm{init} = C/[4a_\mathrm{init}H(a_\mathrm{init})]$, where C is some constant (see for details [7]). The solutions of the equation for matter and dark energy as test components give the asymptotic values for superhorizon perturbations at a_init: $\tilde{\delta}_\mathrm{m}^\mathrm{init} = 3C/4$, $\tilde{v}_\mathrm{m}^\mathrm{init} = C/[4a_\mathrm{init}H(a_\mathrm{init})]$, $\tilde{\delta}_\mathrm{DE}^\mathrm{init} = 3(1+w)C/4$, $\tilde{v}_\mathrm{DE}^\mathrm{init} = C/[4a_\mathrm{init}H(a_\mathrm{init})]$. These relations contain only a single constant C, the value of which specifies the initial amplitudes of perturbations in all components. Below we put $C = 2.6 \cdot 10^{-5}$ for cluster scales $k = 2\,\mathrm{Mpc}^{-1}$ and $C = 6.5 \cdot 10^{-5}$ for supercluster scales $k = 0.2\,\mathrm{Mpc}^{-1}$ at time when $a = 10^{-10}$.

RESULTS AND DISCUSSION

In this section we present the results of numerical integrations of the described system of equations. To achieve this we have designed a Fortran-95 tool based on open `dverk.f` package for ODEs, in which time-dependant functions are replaced by scale-factor dependant functions, since $a(t)$ is well-defined for given cosmology. The cosmological parameters are taken as follows $\Omega_\mathrm{r} = 4.2 \cdot 10^{-5}$, $\Omega_\mathrm{DE} = 0.7$, $\Omega_\mathrm{m} = 1 - \Omega_\mathrm{r} - \Omega_\mathrm{DE}$, $H_0 = 70\,\mathrm{km/s/Mpc}$.

The results of numerical integration of the system of equations for time evolution of amplitudes of dark energy and dark matter density, velocity perturbations, and potential for different scales and parameters of dark energy, are presented in Figs. 1a)–f). In these figures, the solid lines correspond to dark matter perturbations and dashed lines correspond to dark energy perturbations. The dotted lines at all panels show the predictions of linear theory. For such perturbations non-linearity becomes noticeable already at $a \sim 0.1$.

It should be noted that in all cases and for all sets of parameters analysed here, dark energy is sub-dominant: the density of its perturbations is several orders of magnitude lower then that of dark matter. The reason for this is the large absolute value of pressure which keeps it from collapsing. The high effective speed of sound (including speed of pressure perturbation distribution) forces dark energy to oscillate after the perturbation enters the horizon (Figs. 1a) and d)). Both velocity and density perturbation of dark energy are lower throughout the entire course of evolution, and the difference is greater for smaller scales, as can be seen comparing Fig. 1a), b) and Fig. 1d), e).

Another interesting observation is how behaviour of dark energy depends on its parameters. From the Figs. 1a) and c) it is apparent that the EoS parameter w_DE slightly changes the character of evolution of inhomogeneities, while the initial amplitude of density perturbation is $\sim (1 + w_\mathrm{DE})$. At the same time, the effective speed of sound c_s actually defines how fast the perturbation will grow. Comparing Figs. 1a) and b), or d), e) and f) we see that the closer c_s^2 is to zero, the higher amplitude of perturbations is achieved at the final stages of evolution. With lowering of effective speed of sound, the pressure gradient of dark energy lowers the counteraction to gravity. The oscillations become smaller and even disappear at very low values of c_s, while the perturbations grow at a higher rate, and the time evolution of the velocity and density become more and more similar to that of dark matter. This fact is in good accordance with the nature of dependence on effective speed of sound in the process of accretion of dark energy on a compact object, which was studied recently [8]. As can be seen from Fig. 1f), when $c_s = 0$ the dark energy and dark matter always have the same velocity, the gap between their densities is constant, and the velocity factor in the continuity equation for dark energy is $\sim (1 + w_\mathrm{DE})$ (3).

These two properties, together with the observation that dark energy background density is lower than that of components in the past, imply that dark energy could significantly influence the process of dark matter halo collapse, but only under the condition of a very low value of effective speed of sound. The current cosmological observational data do not exclude such models of dark energy ([13]).

It should also be noted that predictions of density and velocity of the linear (doted lines) and non-linear theories at a smaller scale become distinguishable approximately at the same time (scale factor) as for a larger scale. However, for predictions of potential, both linear and non-linear theories diverge only at late times for a smaller scale. The reason for this is, different initial perturbations were taken for these scales, in order to have approximately the same density perturbation at $a = 1$. For the same initial amplitudes the perturbations of smaller scales would collapse first.

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APPENDIX.

NEWTONIAN APPROXIMATION

OF EQUATIONS FOR EVOLUTION

OF DARK ENERGY PERTURBATIONS

Conservation equation $T^k_{0;k}$ transforms to classic continuity equation for dark energy:

$$\dot{\delta}_{DE} + 3\frac{\dot{a}}{a}(c_s^2 - w_{DE})\delta_{DE} + a^{-1}\left\{(1 + c_s^2)\delta'_{DE}v + \right.$$

$$\left. + \left(1 + w + (1 + c_s^2)\delta\right)\left(v' + \frac{2v}{r}\right)\right\} = 0.$$

Combination of conservation equations $T^k_{1;k0}$ and $T^k_{0;k}$ transform to classic Euler equation for dark energy:

$$\dot{v} + \frac{v}{a}\left(v' + \frac{2v}{r}\right) +$$

$$+ v\frac{\dot{a}}{a}\frac{(1 - 3c_s^2)(1 + w_{DE}) + (1 + c_s^2)(1 - 3w)\delta_{DE}}{1 + w_{DE} + (1 + c_s^2)\delta_{DE}} +$$

$$+ \frac{\nu'}{2a} + \frac{c_s^2\delta'_{DE}}{a(1 + w_{DE} + (1 + c_s^2)\delta_{DE})} = 0.$$

If we consider the last two equations for dark matter ($c_s^2 = w_{DE} = 0$), we will see that they coincide with well-known hydrodynamic equation of collapse [12] (equation 9.17)

The non-relativistic approximation of Einstein equation (5) gives the Poisson equation for metric function ν in the coordinates of FRW frame

$$\Delta\nu = 8\pi G a^2\rho_{cr}^0\Big[\Omega_m a^{-3}\delta_m +$$

$$+ \Omega_{de}(1 + 3c_s^2)a^{-3(1+w_{de})}\delta_{de} + 2\Omega_r a^{-4}\delta_r\Big].$$

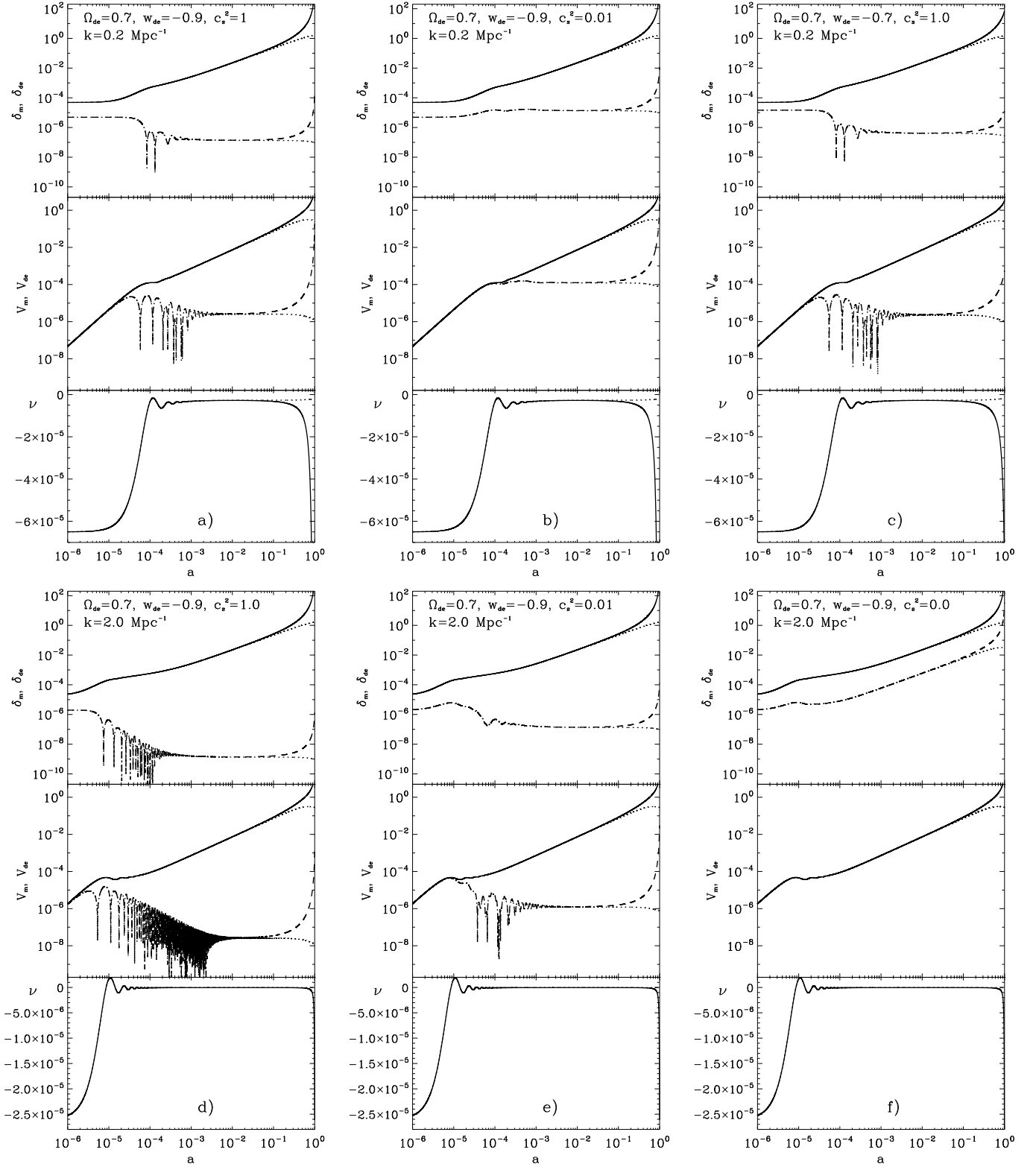The Newtonian gravitational potential $\Phi = \nu/2$.

Fig. 1: Evolution of the amplitudes of density perturbations of matter $\tilde{\delta}_{\rm m}$ and dark energy $\tilde{\delta}_{\rm DE}$ (top panels), velocity of matter $V_{\rm m}$ and dark energy $V_{\rm DE}$ in the units of Hubble one (middle panels) and gravitational potential $\tilde{\nu}$ (bottom panels) at the central part of spherical halo of cluster scale which is collapsing now. In the top and middle panels solid lines corresponds to matter, dashed lines to dark energy and dotted lines at all panels show the prediction of linear theory.